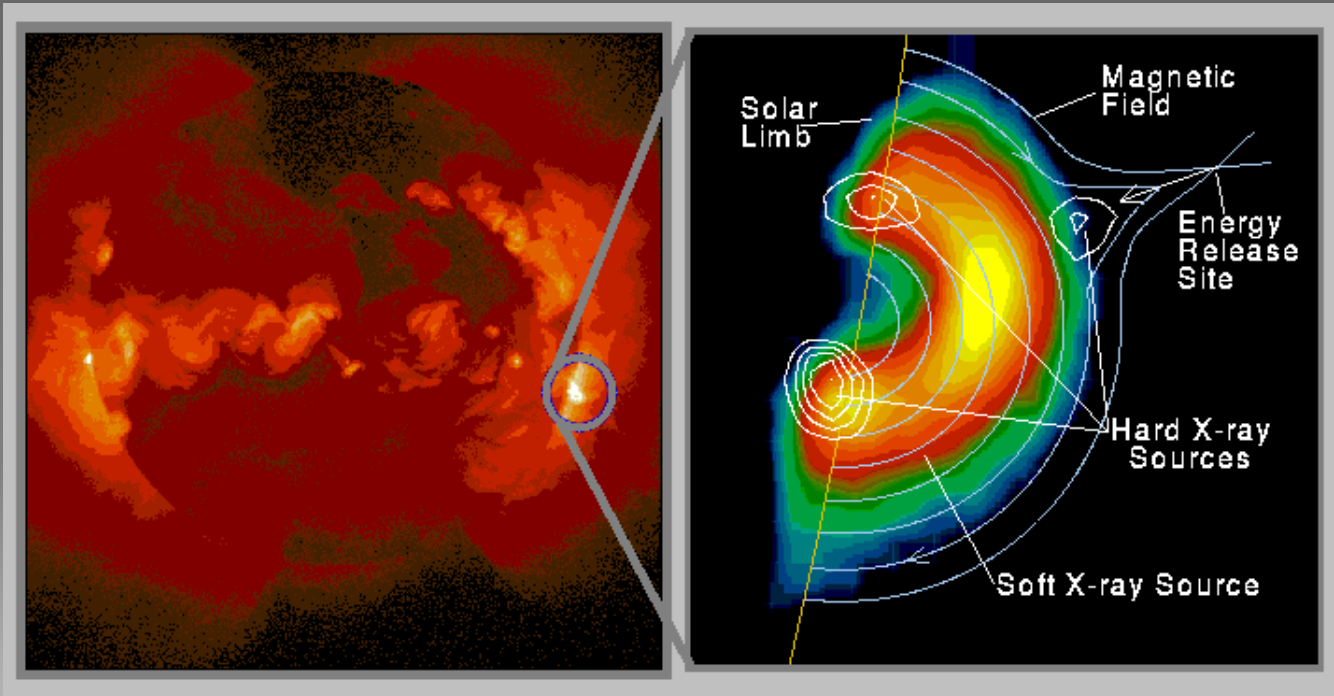


PARTICLE ACCELERATION IN TURBULENT SOLAR ACTIVE REGIONS



Yohkoh X-ray Image of a Solar Flare, Combined Image in Soft X-rays (left) and Soft X-rays with Hard X-ray Contour (right. Jan 13, 1992.)

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M.K. Georgoulis, RCAAM of the Academy of Athens, 4 Soranou Efesiou Street, 11527 Athens, Greece

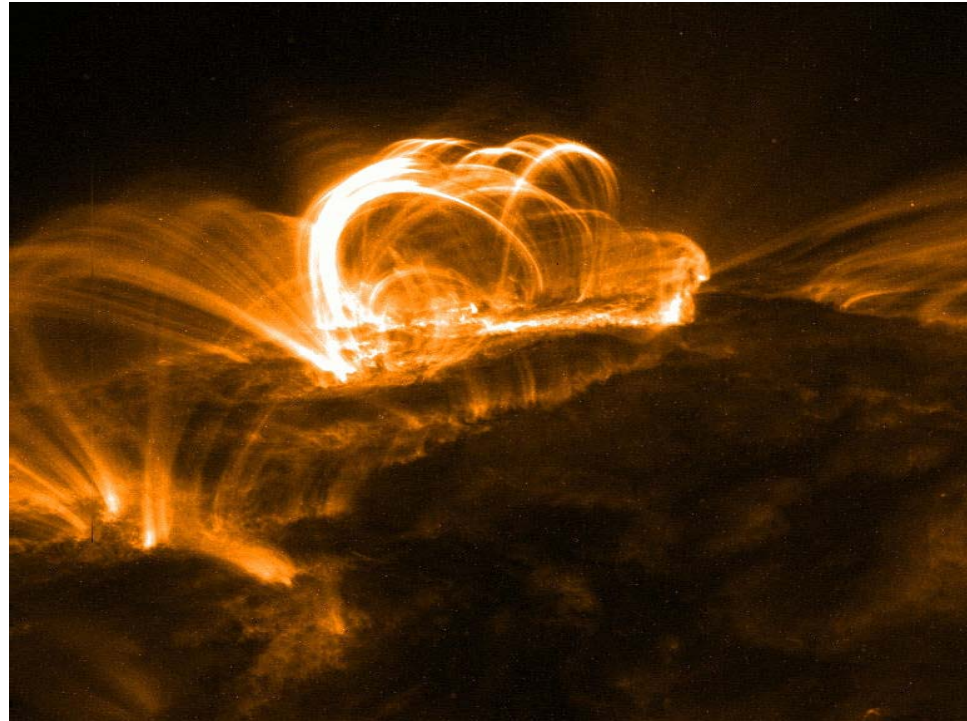
A. Anastasiadis, National Observatory of Athens, IAASARS, GR-15236 Penteli, Greece

SOLAR FLARES AND RECONNECTION SITES

*X-ray & Microwave
Electron associated emission
Gamma-ray
Ion associated emission*

Solar Active Region formation:

*Large-scale
magnetic topology
driven by
strong turbulent flows
evolves forming
magnetic discontinuities
which dissipate
the excess magnetic energy
via reconnection*



2005, TRACE (X-ray), NASA/LMSAL

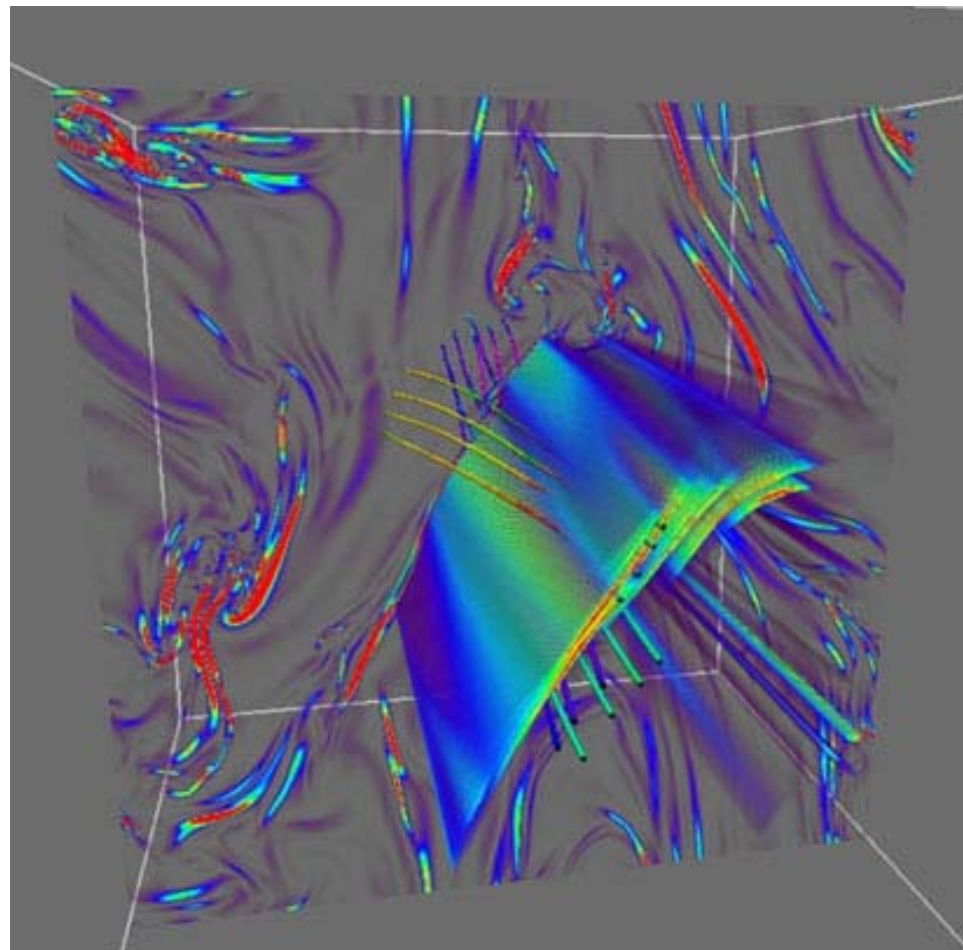
*Current Sheets
Magnetic energy
dissipation regions host
strong electric fields
(also shocks & turbulent flows)
Acceleration sites*

RECONNECTION SITES' STRUCTURE

- *Single elongated current sheet*
(mostly large flare attribute)
- *Multiple small-scale current sheets*
(typical flares)

The *hierarchical fragmentation* of UCSs has been demonstrated with

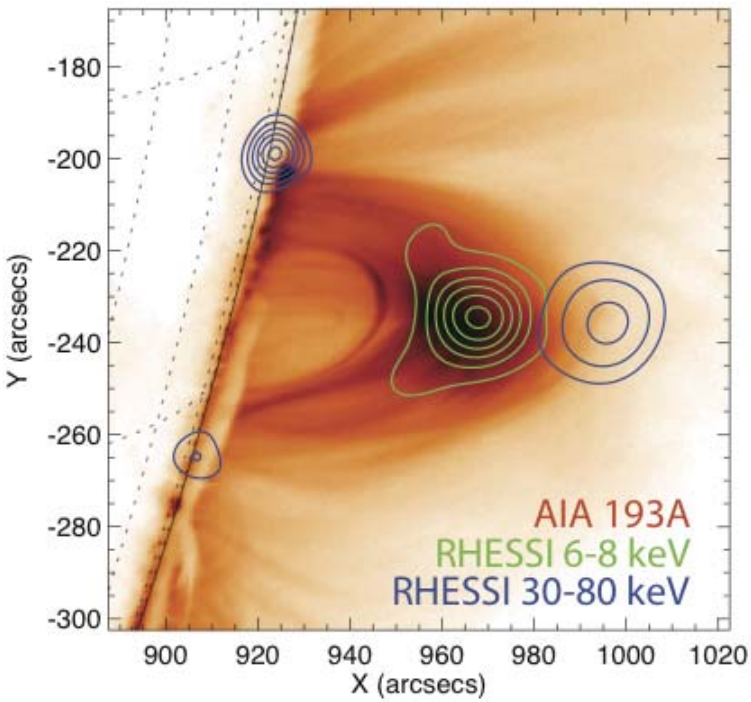
- *MHD*
- *SOC* (Self – Organized Criticality)



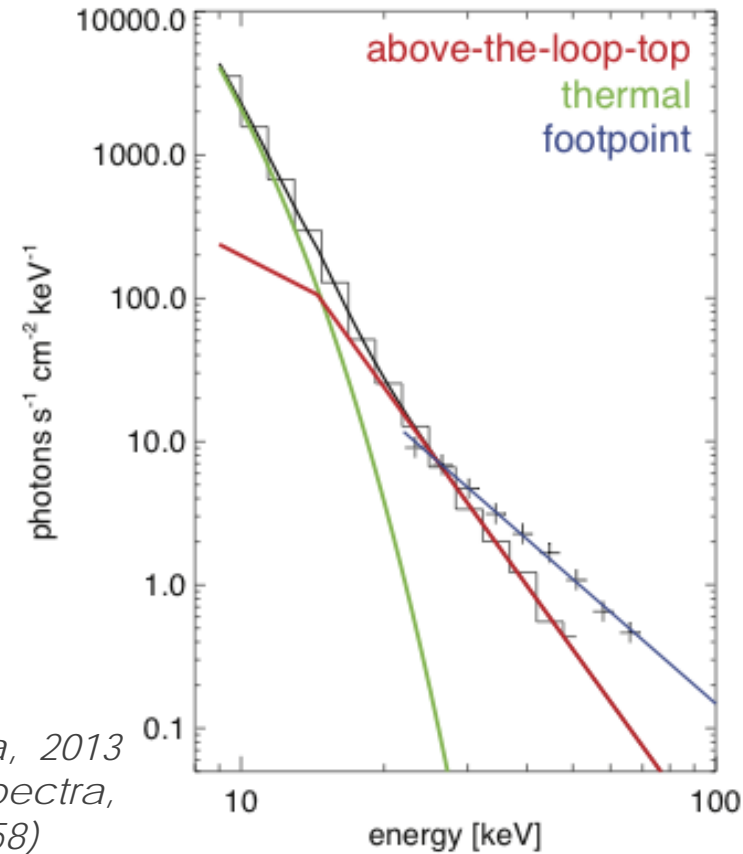
Nordlund and Galsgaard, 2012
(MHD high – resolution braiding simulation)

PARTICLE ACCELERATION IN SOLAR FLARES

One of the requirements for an efficient particle acceleration mechanism is to be able to reproduce the observational emission spectra that is, yield kinetic energy distributions with certain properties



*Krucker and Battaglia, 2013
(RHESSI hard X-ray spectra,
SOL 2012 - 07 - 19 T05:58)*



This task has so far been performed using

- *MHD* (particle trajectories in frozen structures)
- *Kinetic* (direct energy distributions)

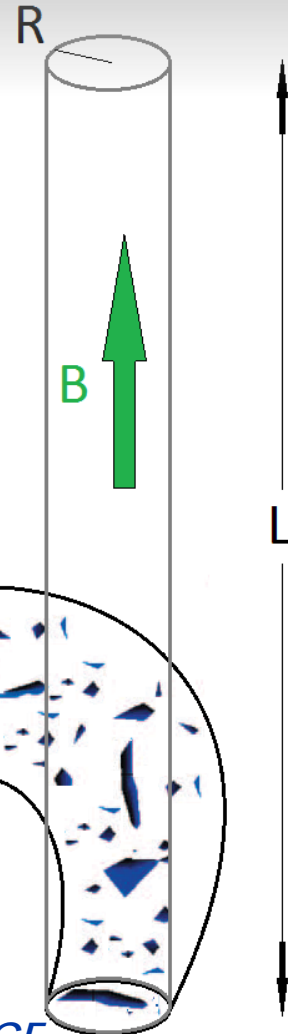
models

PARTICLE ACCELERATION IN TURBULENT SOLAR ACTIVE REGIONS

Toutountzi et al. *Hel.A.S. 11th Conference, 2013*

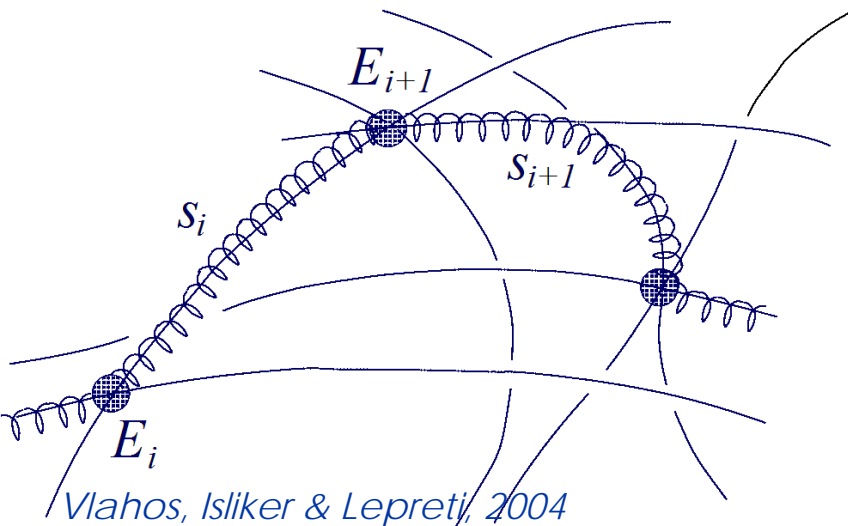
SETTING A.

- Cylindrical volume simulating a **flux tube**
- Initially occupied by **maxwellian plasma**
- Of **uniform density** and temperature $T = 10^6 K$
- **Uniform and stationary** magnetic field of magnitude $B = 10^2 G$ directed upwards along the axis
- **Test-electrons** randomly distributed with randomly directed velocities
- A **particle code** traces their **guiding center** trajectories UCS



SETTING B.

Statistical description of the *electric field's* fractal structure



3 PDFs

$$P_1 = P_1(s)$$

FREE TRAVEL DISTANCE

$$P_2 = P_2(E)$$

ELECTRIC FIELD STRENGTH

$$P_3 = P_3(d)$$

ACCELERATION LENGTH

Vlahos, Isliker & Lepreti, 2004

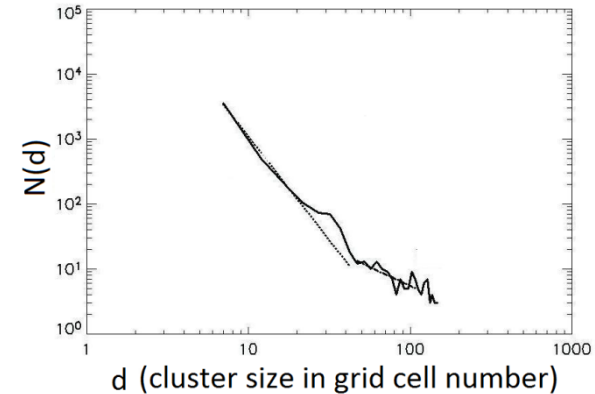
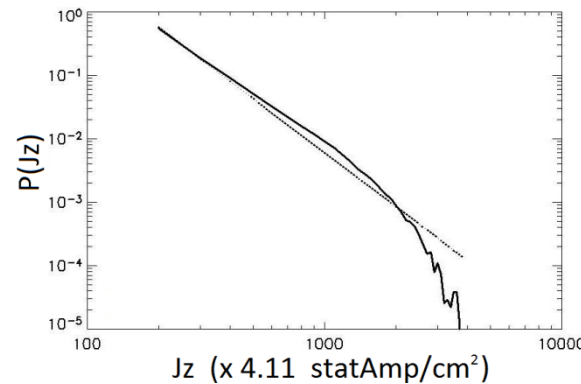
A STATISTICAL STUDY OF CURRENT-SHEET FORMATION ABOVE SOLAR ACTIVE REGIONS BASED ON SELF-ORGANIZED CRITICALITY

Dimitropoulou et al. Hel.A.S. 11th Conference, 2013

BASIC PROCESS STEPS

- Pre - process 2D vector magnetograms from IVM
- 3D magnetic field extrapolation
- Unstable site identification
- Force the system into the SOC state

The formation of unstable clusters is triggered by an avalanche which, as a whole, simulates a solar flare



FREE TRAVEL DISTANCE

$$P_1(s) = \begin{cases} A_1 s^{D_F-3} & D_F \leq 2 \\ A_2 \exp[a_2 s^{D_F-2}] s^{D_F-3} & D_F > 2 \end{cases}$$

Isliker & Vlahos, 2003

ELECTRIC FIELD STRENGTH

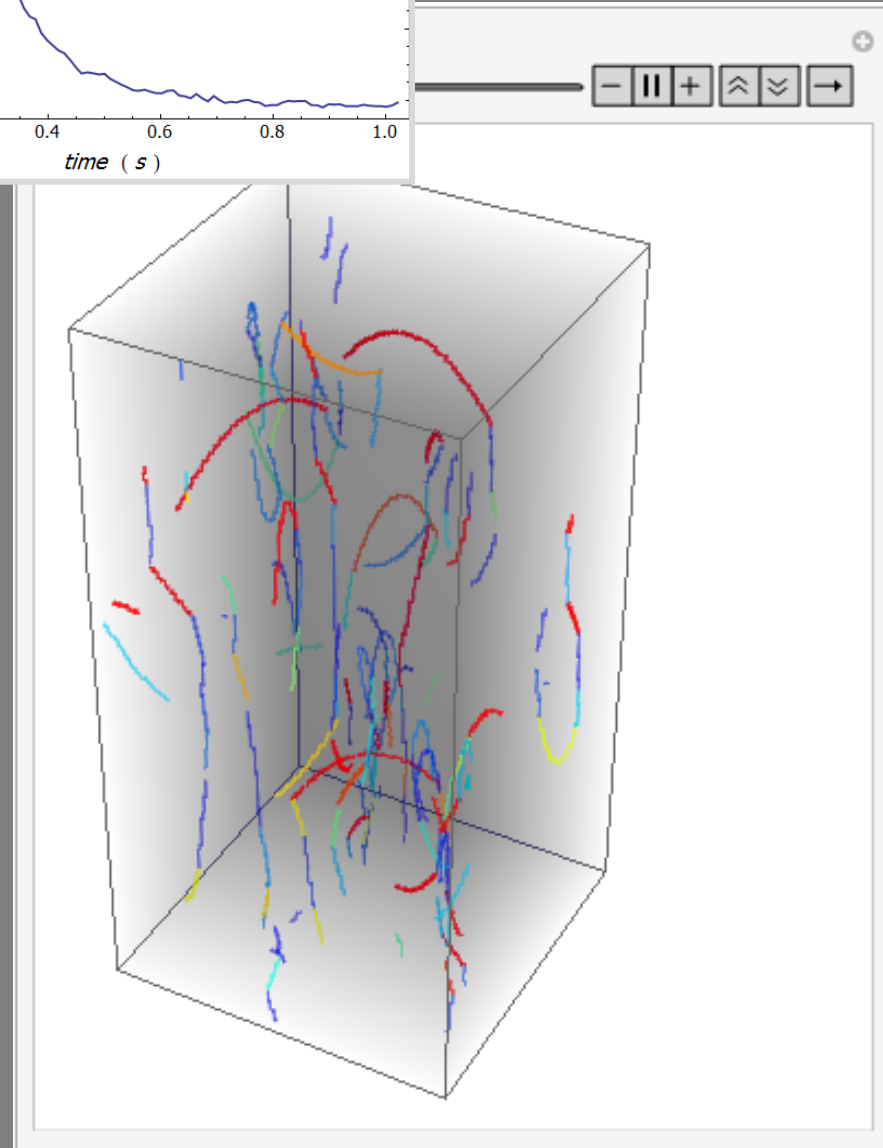
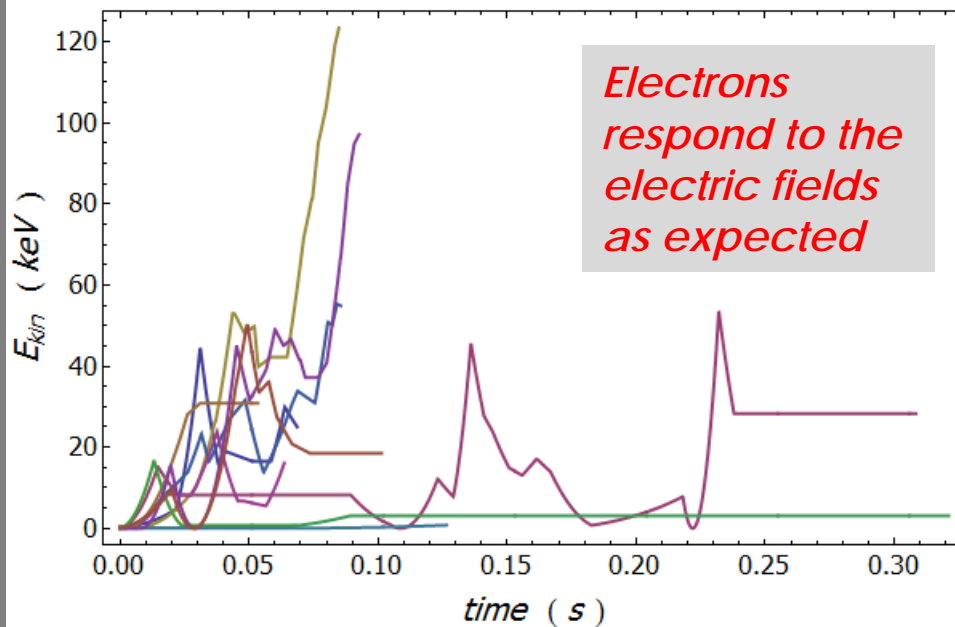
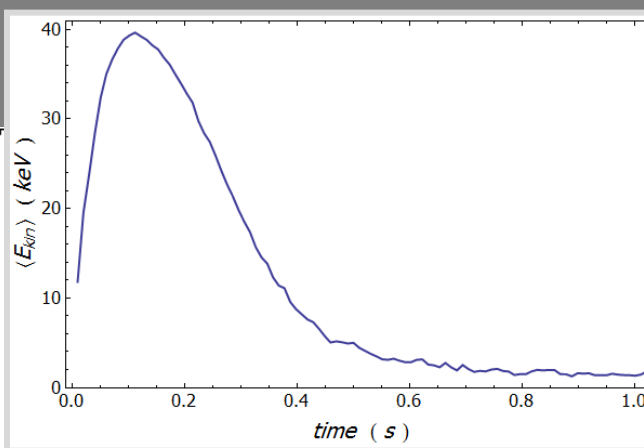
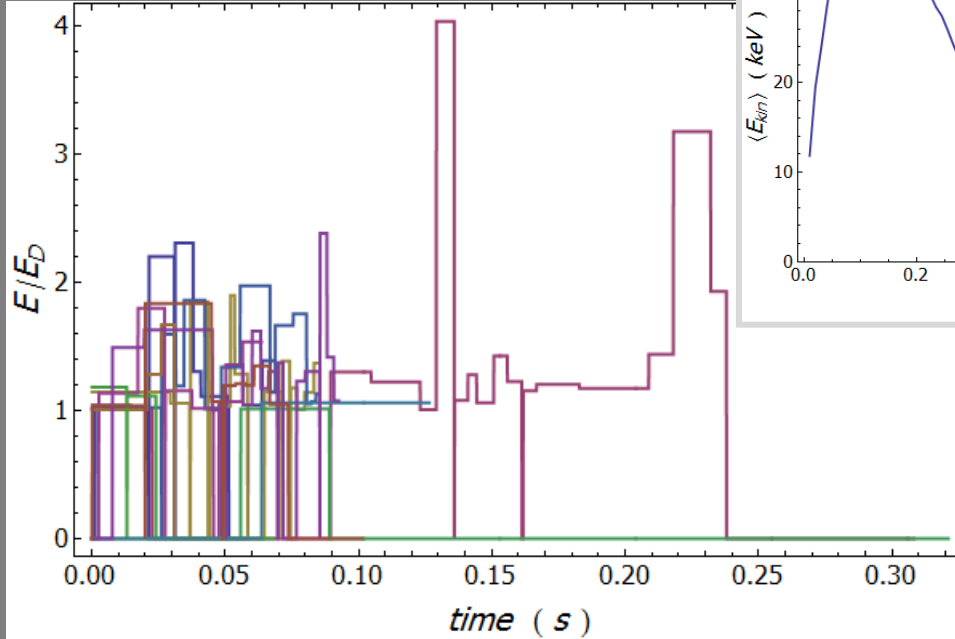
$$P_2(E) = B E^{-b}$$

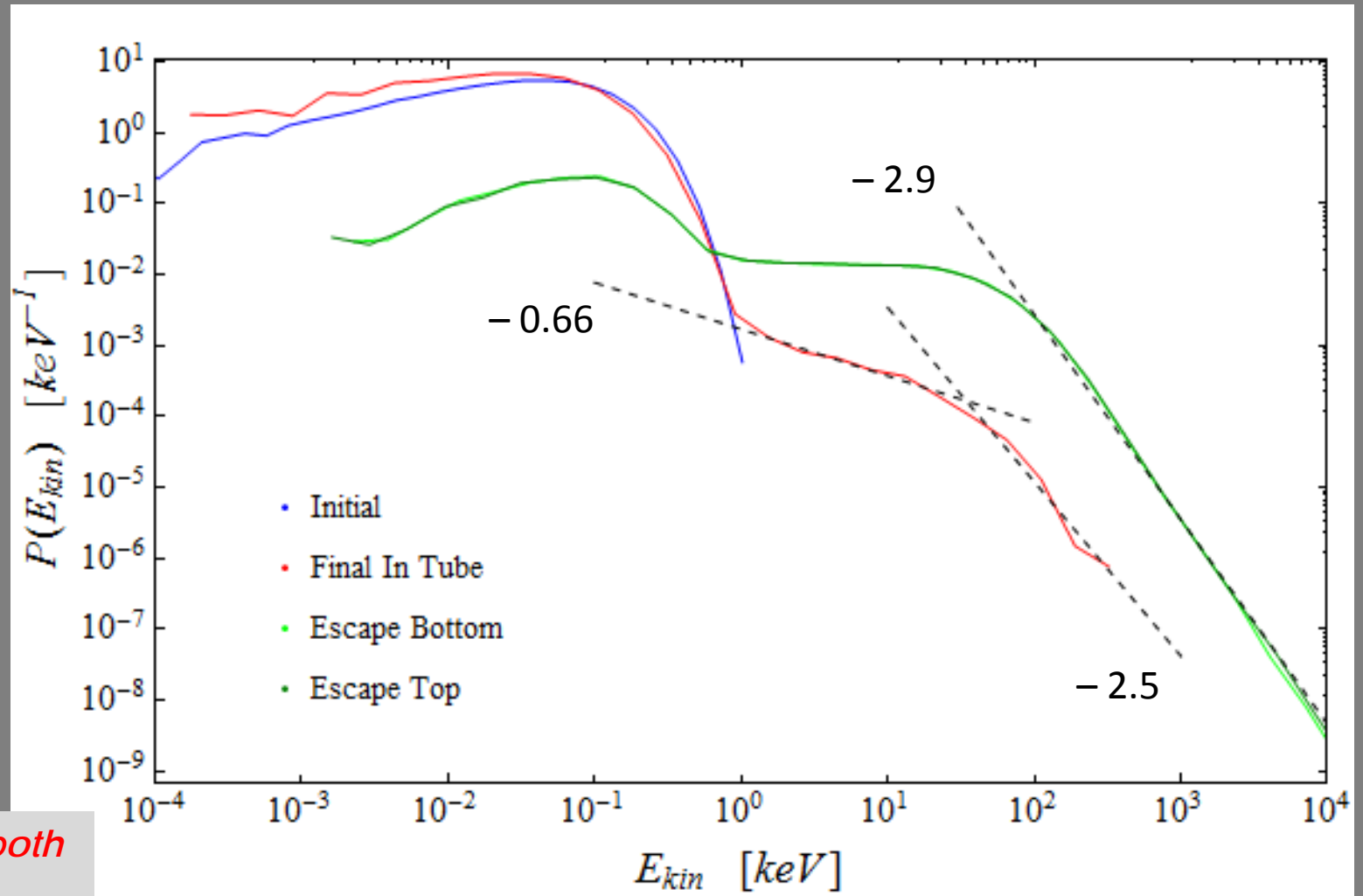
ACCELERATION LENGTH

$$P_3(d) = \begin{cases} C_1 d^{-c_1} & d \leq d_{th} \\ C_2 d^{-c_2} & d > d_{th} \end{cases}$$

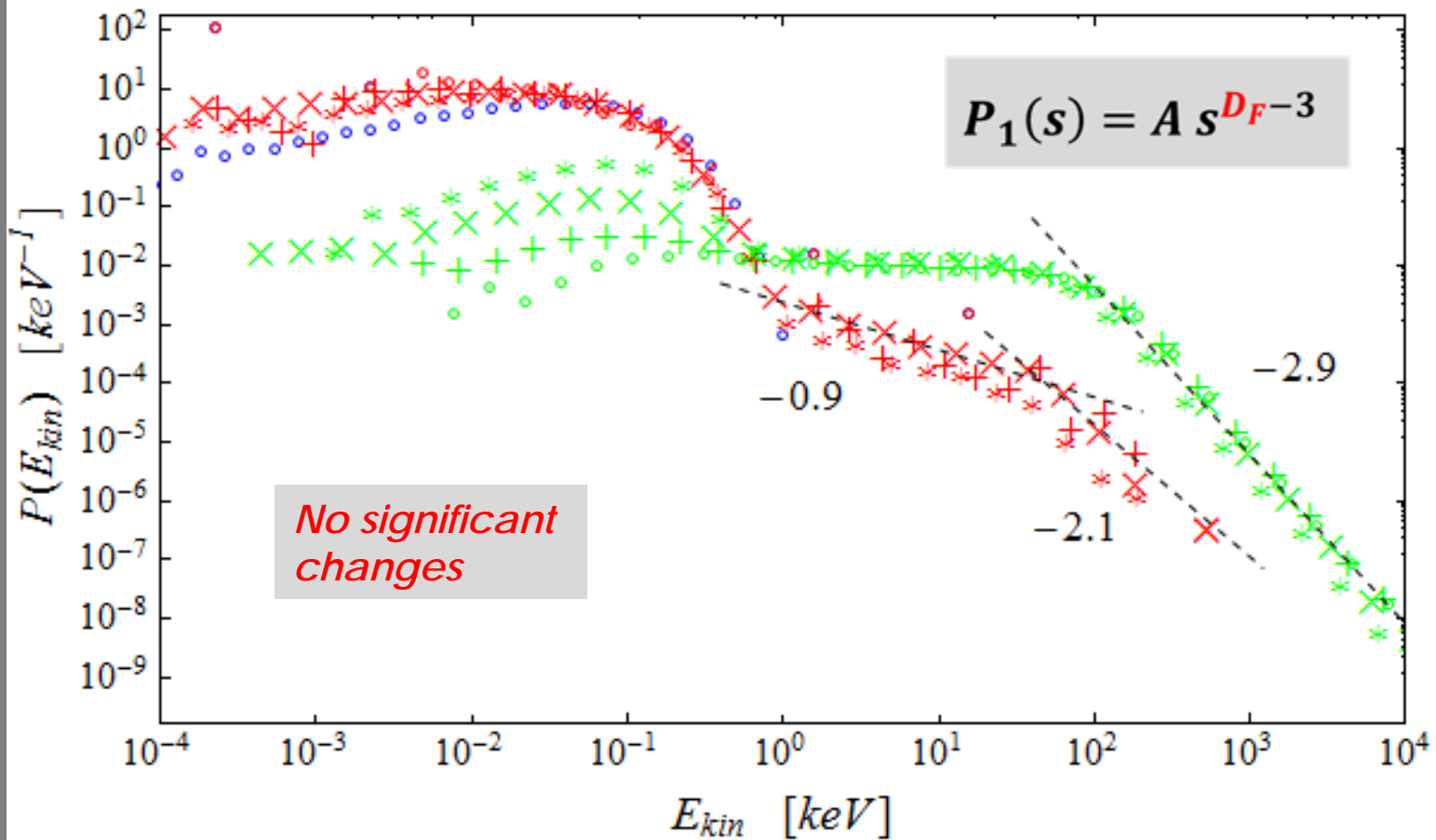
RESULTS 1/5

A FIRST LOOK ON
HOW THINGS WORK





Note that both escaping populations have the same distributions

 **$D_F = 1.3$**

- Initial
- Final In Tube
- Escape

 $D_F = 1.6$

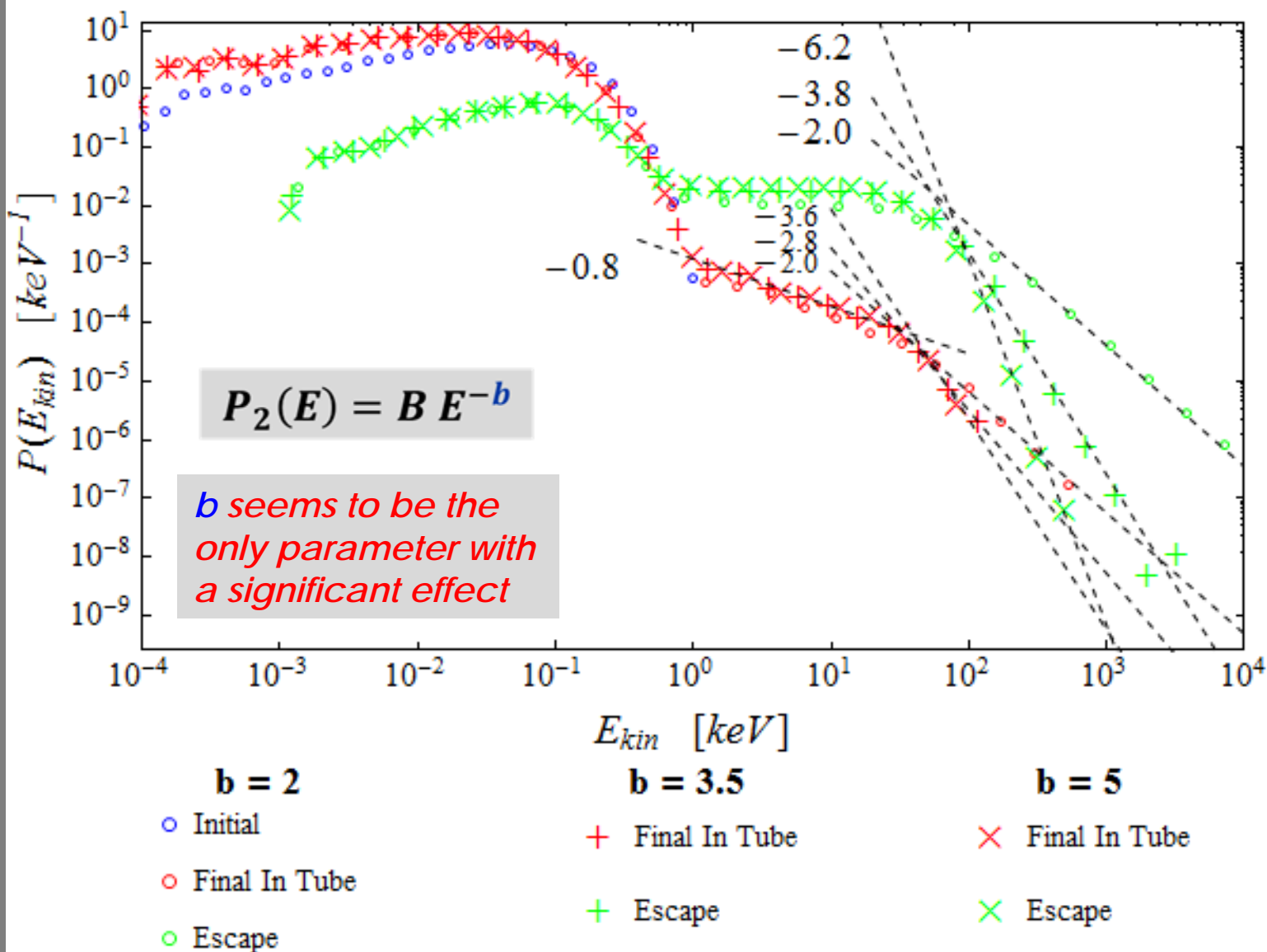
- + Final In Tube
- + Escape

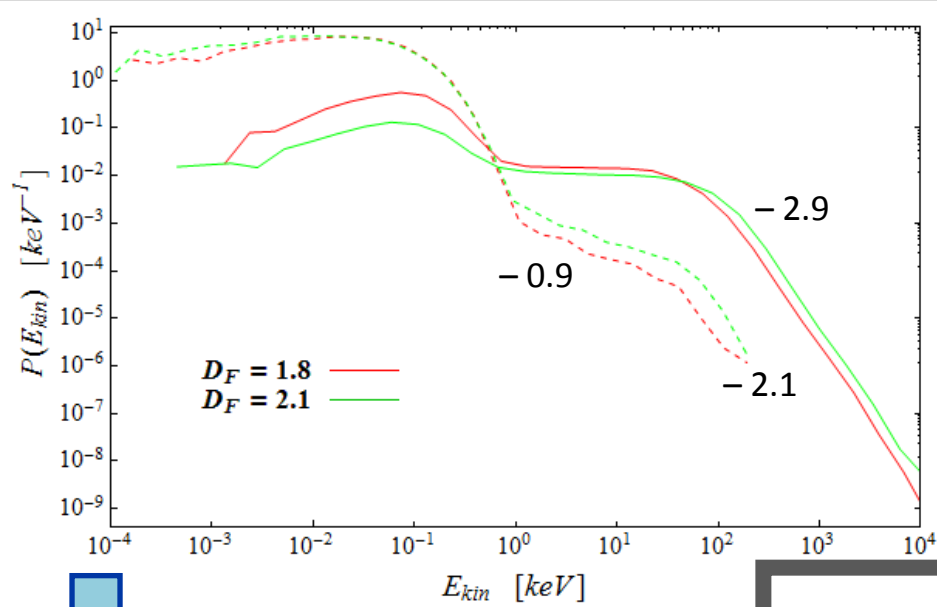
 $D_F = 1.9$

- * Final In Tube
- * Escape

 $D_F = 2.1$

- × Final In Tube
- × Escape



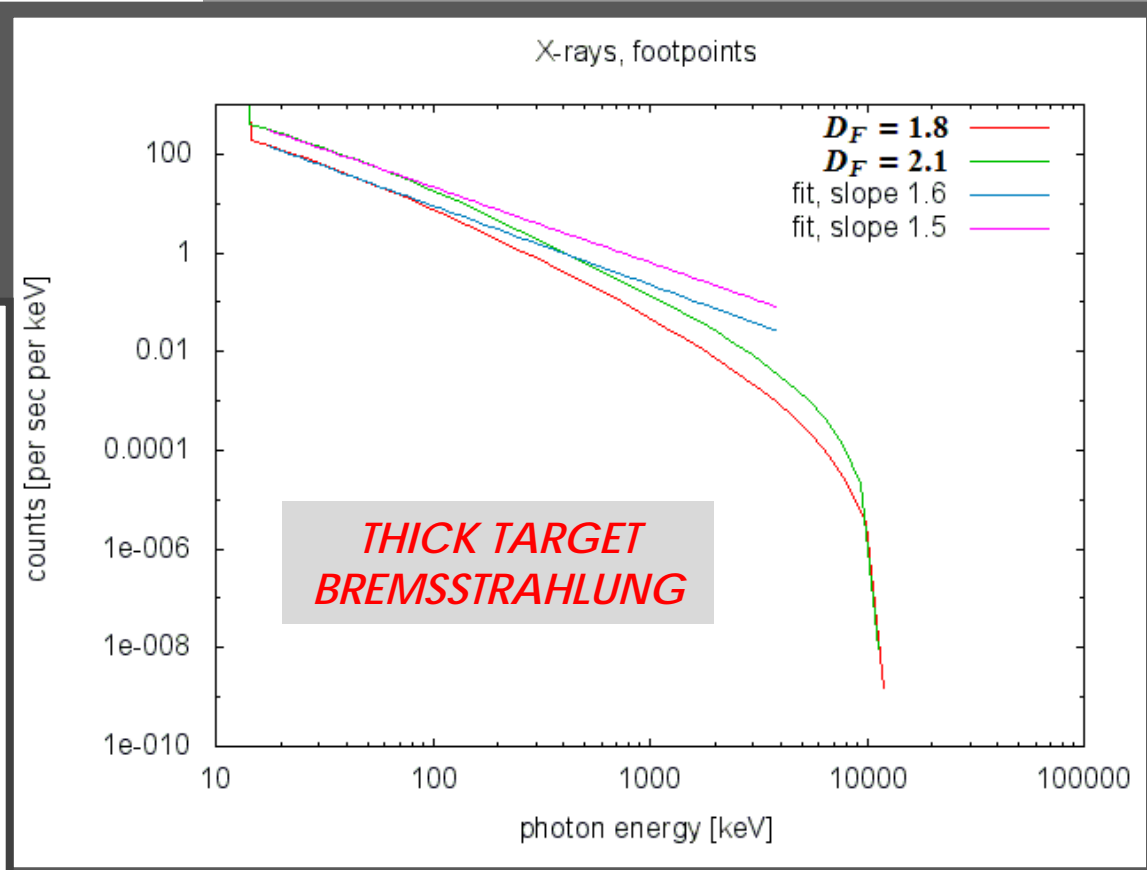
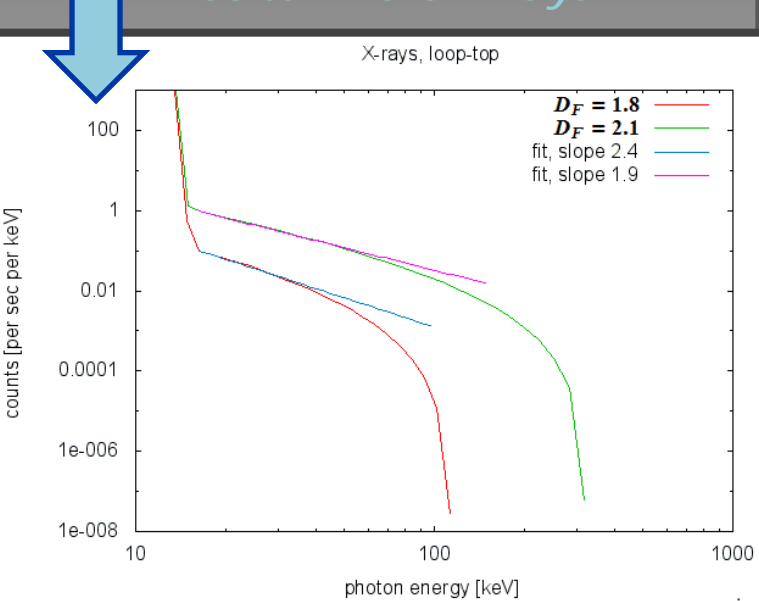


*kinetic energy distribution
of a footpoint source
emitting
Hard X-rays*



**HARD X-RAY
EMISSION SPECTRA**

*kinetic energy distribution
of a loop-top source
emitting
"Softer" Hard X-rays*



SUMMARY / CONCLUSIONS

- *A non-collisional particle code*
 - *traces 500,000 test – electrons'*
 - *guiding center trajectories*
 - *for a time $t = 2$ sec inside a simulated flux – tube with*
 - *a uniform & stationary magnetic field and*
 - *a fragmented electric field*
 - *collecting the particles' energies at various instants*
-
- *Electrons seem to **respond** to the imposed fields as expected*
 - ***Acceleration** actually does take place*
 - *A **parametric study** is performed, revealing that while the*
 - *total time of integration*
 - *maximum electric field strength*
 - *fractality of the electric field structure**do not affect the particle acceleration,*
 - *the **power-law index of the electric field** appears to have*
a non negligible effect
 - *The results concerning the simulated **hard X-ray emission spectra**, even though **preliminary**, they do reveal a power law in the range 20KeV-1MeV for the footpoints and 20 -100KeV for the trapped particles*